

VLT/SPHERE exploration of the young multiplanetary system PDS70*

D. Mesa¹, M. Keppler², F. Cantalloube², L. Rodet³, B. Charnay⁴, R. Gratton¹, M. Langlois^{5,6}, A. Boccaletti⁴, M. Bonnefoy³, A. Vigan⁶, O. Flasseur⁷, J. Bae⁸, M. Benisty^{3,9}, G. Chauvin^{3,9}, J. de Boer¹⁰, S. Desidera¹, T. Henning², A.-M. Lagrange³, M. Meyer¹¹, J. Milli¹², A. Müller², B. Pairet¹³, A. Zurlo^{14,15,6}, S. Antonucci¹⁶, J.-L. Baudino¹⁷, S. Brown Sevilla², E. Cascone¹⁸, A. Cheetham¹⁹, R.U. Claudi¹, P. Delorme³, V. D'Orazi¹, M. Feldt², J. Hagelberg¹⁹, M. Janson²⁰, Q. Kral⁴, E. Lagadec²¹, C. Lazzoni¹, R. Ligi²², A.-L. Maire^{2,23}, P. Martinez²¹, F. Menard³, N. Meunier³, C. Perrot^{4,24,25}, S. Petrus³, C. Pinte^{26,3}, E.L. Rickman¹⁹, S. Rochat³, D. Rouan⁴, M. Samland^{2,20}, J.-F. Sauvage^{27,6}, T. Schmidt^{4,28}, S. Udry¹⁹, L. Weber¹⁹, F. Wildi¹⁹

(Affiliations can be found after the references)

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ABSTRACT

Context. PDS 70 is a young (5.4 Myr), nearby (~113 pc) star hosting a known transition disk with a large gap. Recent observations with SPHERE and NACO in the near-infrared (NIR) allowed us to detect a planetary mass companion, PDS 70 b, within the disk cavity. Moreover, observations in H_α with MagAO and MUSE revealed emission associated to PDS 70 b and to another new companion candidate, PDS 70 c, at a larger separation from the star. PDS 70 is the only multiple planetary system at its formation stage detected so far through direct imaging.

Aims. Our aim is to confirm the discovery of the second planet PDS 70 c using SPHERE at VLT, to further characterize its physical properties, and search for additional point sources in this young planetary system.

Methods. We re-analyzed archival SPHERE NIR observations and obtained new data in Y, J, H and K spectral bands for a total of four different epochs. The data were reduced using the data reduction and handling pipeline and the SPHERE data center. We then applied custom routines (e.g. ANDROMEDA and PACO) to subtract the starlight.

Results. We re-detect both PDS 70 b and c and confirm that PDS 70 c is gravitationally bound to the star. We estimate this second planet to be less massive than $5 M_{\text{Jup}}$ and with a T_{eff} around 900 K. Also, it has a low gravity with $\log g$ between 3.0 and 3.5 dex. In addition, a third object has been identified at short separation (~0.12") from the star and gravitationally bound to the star. Its spectrum is however very blue, so that we are probably seeing stellar light reflected by dust and our analysis seems to demonstrate that it is a feature of the inner disk. We, however, cannot completely exclude the possibility that it is a planetary mass object enshrouded by a dust envelope. In this latter case, its mass should be of the order of few tens of M_\oplus . Moreover, we propose a possible structure for the planetary system based on our data that, however, cannot be stable on a long timescale.

Key words. Instrumentation: spectrographs - Methods: data analysis - Techniques: imaging spectroscopy - Stars: planetary systems, Stars: individual: PDS 70

1. Introduction

PDS 70 is a 5.4 ± 1.0 Myr (Müller et al. 2018) K7 pre-main sequence star that is part of the Upper Centaurus-Lupus group (Pecaut & Mamajek 2016) at a distance of 113.43 ± 0.52 pc (Gaia Collaboration et al. 2016, 2018). In the course of the VLT/SPHERE SHINE survey (Chauvin et al. 2017), Keppler et al. (2018) discovered a planetary mass companion, PDS 70 b, located in the transition disk surrounding this young star.

The presence of a circumstellar disk was first inferred by Metchev et al. (2004) due to the detection of a strong mid-infrared excess and of a strong emission at millimeter wavelengths. The disk was first resolved in K_s band by Riaud et al. (2006) using VLT/NACO. The first detection of a gap in the disk was obtained by Hashimoto et al. (2012) exploiting H-band polarized data obtained through the Subaru/HiCIAO instrument. Using the same data, Dong et al. (2012) estimated, assuming a

distance of 140 pc for the system, a gap size of around 65 au in which the dust is depleted by a factor of ~1000 with respect to the outer part of the disk. They also found evidence of an inner disk with dimensions of the order of few au detecting a weak near-IR excess in the spectral energy distribution (SED). Finally, they estimated for the disk a total dust mass of $\sim 10^{-4} M_\odot$. PDS 70 was also observed at millimeter wavelengths by Hashimoto et al. (2015) using the sub-millimeter array (SMA) and, more recently, by Long et al. (2018) using ALMA. The ALMA observations were performed both at 0.87 mm continuum and at HCO^+ and CO gas emission lines. They allowed to define, in the dust continuum, the presence of a radial gap between the inner and the outer disk at 15-60 au of separation, assuming a distance of 140 pc, and to image different substructures as a bridge-like feature and an azimuthal gap in the HCO^+ emission of the disk. Further observations with ALMA, in the continuum and CO, at higher angular resolution were presented by Keppler et al. (2019) and showed evidence of a depletion of emission in the CO integrated intensity centered on the separation of PDS 70 b, while the continuum peak is located at 74 au. Moreover, they found, through hydrodynamical modelling of the

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gas kinematics, that the presence of an additional low-mass companion further out than the orbit of PDS 70 b may be required to account for the large gap width.

PDS 70 b was detected into the disk gap using NICI, NACO and VLT/SPHERE data in the near-infrared (NIR) by [Keppler et al. \(2018\)](#). Its colors are very red and through the photometry they were able to estimate a mass of the order of $5\text{--}9 M_{\text{Jup}}$ assuming for the companion the same age of the system. The outer disk, as seen at NIR wavelengths, has a radius of about 54 au assuming the updated value of 113.43 pc for the system. Moreover, for the first time they were able to detect scattered light from the inner disk determining a radius of less than 17 au, consistent with ALMA values. Exploiting new SPHERE observations [Müller et al. \(2018\)](#) performed a characterization of the orbital properties of PDS 70 b constraining its semi-major axis around 22 au corresponding to an orbital period of ~ 118 years. Using the planet infrared spectrum and atmospheric models, they found for the companion a T_{eff} between 1000 and 1600 K, low surface gravity (< 3.5 dex) and an unusually large radius between 1.4 and $3.7 R_{\text{Jup}}$. The latter could be a hint of the presence of circumplanetary material. Recent observations in H and K band with VLT/SINFONI may indicate the presence of a circumplanetary disk around PDS 70 b to explain its very red spectrum ([Christians et al. 2019a,b](#)).

Following the PDS 70 b discovery, observations at $H\alpha$ wavelength with the Magellan Adaptive Optics (MagAO) system ([Wagner et al. 2018](#)) revealed a source in $H\alpha$ at the position of the companion. This is a hint that PDS 70 b is still accreting material and is consistent with the presence of a circumplanetary disk. The authors were able to derive a value of $10^{-8\pm 1} M_{\text{Jup}} \text{ yr}^{-1}$ for the mass accretion rate. This result was later independently confirmed with VLT/MUSE observations ([Haffert et al. 2019](#)) who redetected PDS 70 b in $H\alpha$ and derived an accretion rate of $2 \times 10^{-8\pm 0.4} M_{\text{Jup}} \text{ yr}^{-1}$. A new clear point-source emission was detected with a signal-to-noise (S/N) of around 8 at larger separation (~ 35 au) very near to the bright west ring of the disk as seen in projection, hinting for the presence of a second companion in the PDS 70 system (hereinafter PDS 70 c). The authors were able to estimate an accretion rate of $10^{-8\pm 0.4} M_{\text{Jup}} \text{ yr}^{-1}$ for this object. The reanalysis of the ALMA data previously used by [Long et al. \(2018\)](#) and by [Keppler et al. \(2019\)](#) allowed [Isella et al. \(2019\)](#) to detect a sub-millimeter continuum emission associated with PDS 70 c interpreted as originating from a dusty circumplanetary disk. They also detected a second compact source at close separation from PDS 70 b speculating that this could come from dust orbiting in the proximity of the planet.

The direct imaging signal of PDS 70 c was marginally identified in SPHERE NIR data in the form of a rather elongated structure but not presented in [Keppler et al. \(2018\)](#) and [Müller et al. \(2018\)](#) because of possible contamination from the disk. Further SPHERE follow-up observations were performed with the aim to provide a robust independent confirmation of the existence of PDS 70 c, to further characterize its nature and, finally, to search for additional point-sources in the system.

In this paper, we present in Section 2 the data that we use for our analysis and the data reduction procedure. In Section 3 we present our results while in Section 4 we discuss them. Finally, in Section 5 we give our conclusions.

2. Observations and data reduction

For the present work we have used both archival and new observations taken with SPHERE ([Beuzit et al. 2019](#)). The archival observations were obtained on the nights of 2015-05-31 and

2018-02-24 and were previously used for the works presented in [Keppler et al. \(2018\)](#) and in [Müller et al. \(2018\)](#). In addition to these data we have also acquired new data on the nights of 2019-03-06 and 2019-04-13. The main characteristics of all these observations are presented in Table 1. The first of these observations was carried out in the IRDIFS mode, that is with IFS ([Claudi et al. 2008](#)) operating in Y and J spectral bands (between 0.95 and $1.35 \mu\text{m}$) and IRDIS ([Dohlen et al. 2008](#)) operating in the H band with the H23 filter pair (wavelength H2= $1.593 \mu\text{m}$; wavelength H3= $1.667 \mu\text{m}$; [Vigan et al. 2010](#)). The remaining observations were performed using the IRDIFS_EXT mode that uses IFS in Y, J and H spectral band (between 0.95 and $1.65 \mu\text{m}$) and IRDIS exploiting the K band with the K12 filter pair (K1= $2.110 \mu\text{m}$ and K2= $2.251 \mu\text{m}$). Due to technical problems, the star was offset by ~ 20 mas in the South-East direction with respect to the coronagraphic mask during the 2019-03-06 observations. In the last observing date we used a coronagraph that was not optimized for the IRDIFS_EXT mode but that had a smaller inner working angle, with a diameter of 145 mas, allowing us to observe possible objects at small separations from the central star. This, however, did not allow to obtain good astrometric and photometric measurements for PDS 70 b and PDS 70 c and for this reason we were not able to use these values for this epoch in the following work. At all the epochs, we obtained frames with satellite spots symmetric with respect to the central star before and after the coronagraphic sequences. This enabled us to determine the position of the star behind the coronagraphic focal plane mask and accurately recenter the data. Furthermore, to be able to correctly calibrate the flux of companions, we have acquired images with the star off-axis. In these cases, the use of appropriate neutral density filter was mandatory to avoid saturation of the detector.

The data were reduced through the SPHERE data center ([DeLorme et al. 2017](#)) applying the appropriate calibrations following the data reduction and handling (DRH; [Pavlov et al. 2008](#)) pipeline. In the IRDIS case, the requested calibrations are the dark and flat-field correction and the definition of the star center. IFS requires, besides to the dark and flat-field corrections, the definition of the position of each spectra on the detector, the wavelength calibration and the application of the instrumental flat. On the pre-reduced data we then applied speckle subtraction algorithms like TLOCI ([Marois et al. 2014](#)) and principal components analysis (PCA; [Soummer et al. 2012](#)) as implemented in the consortium pipeline application called SpeCal (Spectral Calibration; [Galicher et al. 2018](#)) and also described in [Zurlo et al. \(2014\)](#) and in [Mesa et al. \(2015\)](#) for the IFS. Finally, the ANDROMEDA package ([Cantalloube et al. 2015](#)) and the PACO package ([Flasseur et al. 2018](#)) were also applied to all datasets.

3. Results

The final images, obtained using PCA are shown in Figure 1 both for IFS (upper panels) and IRDIS (bottom panels). PDS 70 b is visible in all these images but in the following we will focus on the detection of other point sources.

The re-analysis of these images allowed us to detect PDS 70 c at a S/N ratio of 8.6 and 9.6 at a projected separation of $\sim 0.2''$ in the IRDIS data acquired during the nights of 2018-02-24 and 2019-03-06 (panel *f* and *g* of Figure 1). PDS 70 c is visible at the same epochs in the IFS images (panel *b* and *c* of Figure 1) and, barely, in the IRDIS image obtained from 2019-04-13 data (panel *h* of Figure 1).

A third object is clearly visible at short separation from the star both in the 2018-02-24 and in the 2019-03-06 IFS data

Table 1. List and main characteristics of the SPHERE observations of PDS 70 used for this work. During the observation of 2019-03-06, the coronagraph was incorrectly positioned with an offset of ~ 20 mas in South-East direction.

Date	Obs. mode	Coronagraph	DIMM seeing	τ_0	wind speed	Field rotation	DIT	Total exposure
2015-05-31	IRDIFS	N_ALC_YJH_S	1.20''	1.1 ms	4.55 m/s	50.6°	64 s	4096 s
2018-02-24	IRDIFS_EXT	N_ALC_YJH_S	0.40''	7.0 ms	2.85 m/s	93.4°	96 s	6336 s
2019-03-06	IRDIFS_EXT	N_ALC_K _s	0.39''	8.9 ms	4.45 m/s	56.5°	96 s	4608 s
2019-04-13	IRDIFS_EXT	N_ALC_YJ_S	0.98''	2.5 ms	8.53 m/s	71.4°	96 s	6144 s

(panel *b* and *c* of Figure 1) with a S/N ratio of ~ 5 . As we will point out in the following Sections, even if the planetary nature of this object cannot be fully excluded, our analysis favours that this object is an inner disk feature. For this reason we will hereinafter refer to this detection as a point-like feature (PLF). The source was also retrieved in the IFS 2019-04-13 observation (panel *d* of Figure 1), during which we used a set-up optimized to retrieve objects at very close separation from the star, as explained above. Furthermore, we were able to find the same object at very low S/N in an older observation (2015-05-31, see panel *a* of Figure 1). This allowed us to expand the time range of our observations. For the IRDIS data, this object can only be barely retrieved in the 2018-02-24 data while it is not visible in the remaining epochs. However, given the poor detection, the photometry could not be extracted from these data either using the negative planet method or the ANDROMEDA package. It was instead possible to retrieve a photometry using the PACO package.

We have then combined the images at different epochs to obtain a complete vision of the PDS 70 planetary system and to determine in a more precise way its structure following the method described in Gratton et al. (2019). We have used the IFS images because it was possible to clearly image all of the proposed companions around the star. As a first step we deprojected our data, assuming a disk inclination and a position angle of 49.7° and 158.6° respectively (Hashimoto et al. 2012; Keppler et al. 2019), to be able to see the disk plane face-on. We then rotated the images according to circular Keplerian motion to have the three objects in the same positions at different epochs and sum the images. Given the short total orbital rotation of the companions in the considered time range, the effects of the uncertainties on the stellar and companion masses are not relevant in the final result of this procedure. Finally, we projected back our data to obtain a view of the system as seen on-sky.

The final result of this procedure is shown in Figure 2 where in the left panel we show the on-sky view of the system with the indication of the positions of the three proposed objects. On the right panel, we instead show the deprojected view of the system seen in the disk plane. We also plotted a dashed circle corresponding to the expected position of the inner edge of the outer disk inferred from the sub-millimeter continuum observations (Keppler et al. 2019) to show that it corresponds to the eastern part of the ring visible in the SPHERE image. Furthermore, we note that the position of the western outer ring of the disk, apparently very near to the position of PDS 70 c is due to a projection effect because the bright western part of the disk seen in scattered light is actually the upper layer of a torus at certain height with respect to the disk plane. This is similar to other cases observed with SPHERE like e.g. RX J1615.3-3255 (de Boer et al. 2016) and HD 100546 (Sissa et al. 2018).

A complete discussion of the structure of the disk is however beyond the scope of this paper, which is focused on the confirmation and characterization of PDS 70 c and on the additional point-like feature in this system.

3.1. PDS 70 b

The new SPHERE data did not allow us to update the spectral results given in Müller et al. (2018) for PDS 70 b. We then refer to this paper for any information about the spectral analysis for this object.

We used our recent SPHERE data to update the orbital solutions of PDS 70 b which are in good agreement with what was found by Müller et al. (2018). The results of this analysis are given in the top right panel of Figure 3, in Section 3.4 and in Section 4.3.

3.2. PDS 70 c

The main limitation to the analysis of this object with SPHERE is that it is projected very near to the bright western ring of the disk. This prevented a robust identification in previous analysis as it was regarded as a feature of the disk and, moreover, complicating both the astrometric and photometric extraction. In our new analysis, we were able to obtain these values using the negative planet method (see e.g., Bonnefoy et al. 2011; Zurlo et al. 2014) for the two best observing epochs (2018-02-24 and 2019-03-06) and confirming that it is actually a planet. In Table 2 we list the astrometric values obtained from the IRDIS data for the two available epochs through this approach. The relative positions of PDS 70 c are also shown in the bottom left panel of Figure 3 where they are compared to the positions expected for a background object. This allows us to exclude that this is a background object.

The negative planet method was also used to calculate the photometric values. However, due to the faintness of this object in the Y, J and H spectral bands, we had to perform this method using the average over 5 consecutive spectral channels. This resulted in the fact that we could not derive any photometric value for the first two and the last two IFS channels. The final spectrum is obtained through a weighted average, with the weight considering the different qualities of the observations, of the results obtained for the last three epochs considered here. To check the reliability of this result we have extracted the spectrum also using the ANDROMEDA and the PACO package. ANDROMEDA gave results comparable to those from the first method as shown in Figure 4, apart from the fact that the negative planet method gives negative values in the spectral region around $1.4 \mu\text{m}$. These values are likely related to water telluric absorption affecting both the stellar and the planet flux but we cannot exclude that they are linked to the effect of the data reduction method. The same feature is not present in the ANDROMEDA spectrum. On the other hand, the spectrum extracted with PACO had a detection limit below 3 sigma so that we did not consider it as a reliable detection and we do not use it for this work. Instead, its results for the IRDIS data are in good agreement with those for ANDROMEDA as shown in Figure 4. For all the reason listed above, we will therefore use in the following just the results from the ANDROMEDA package. The result of this procedure is dis-

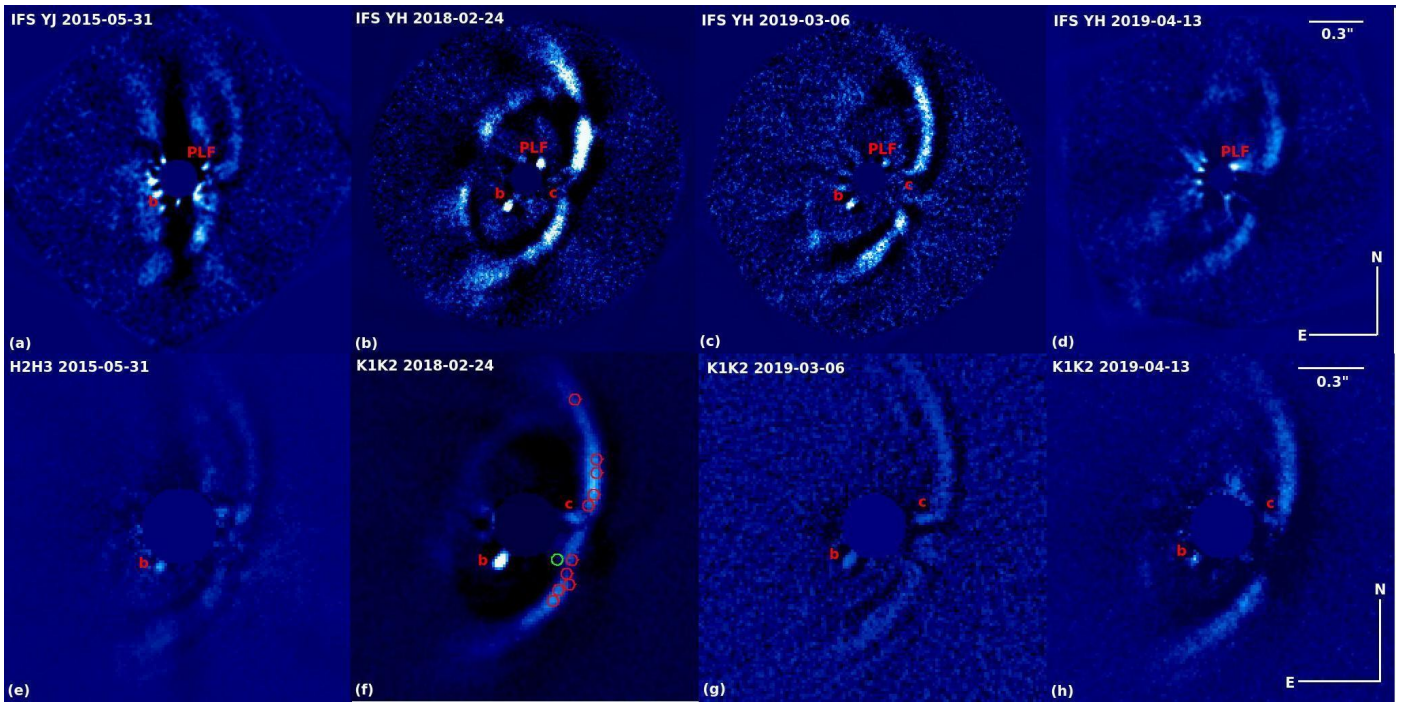


Fig. 1. Final IFS (upper row) and IRDIS (bottom row) images obtained for the following epochs: 2015-05-31 (panel *a* and panel *e*), 2018-02-24 (panel *b* and panel *f*), 2019-03-06 (panel *c* and panel *g*) and 2019-04-13 (panel *d* and panel *h*). The scale and the orientation are the same for all the epochs and are displayed in panel *d* and in panel *h*. In each image we tag with red letters the positions of the identified companions and of the proposed point like feature (PLF). All the images displayed have been obtained using a 5 principal components PCA applied separately on each wavelength apart that for the first IFS image where just 2 principal components were used. In panel *f* we overplot ten red circles to indicate the positions used to extract the disk spectra and one green circle for the position of the simulated planet as described in more details in Section 4.1.

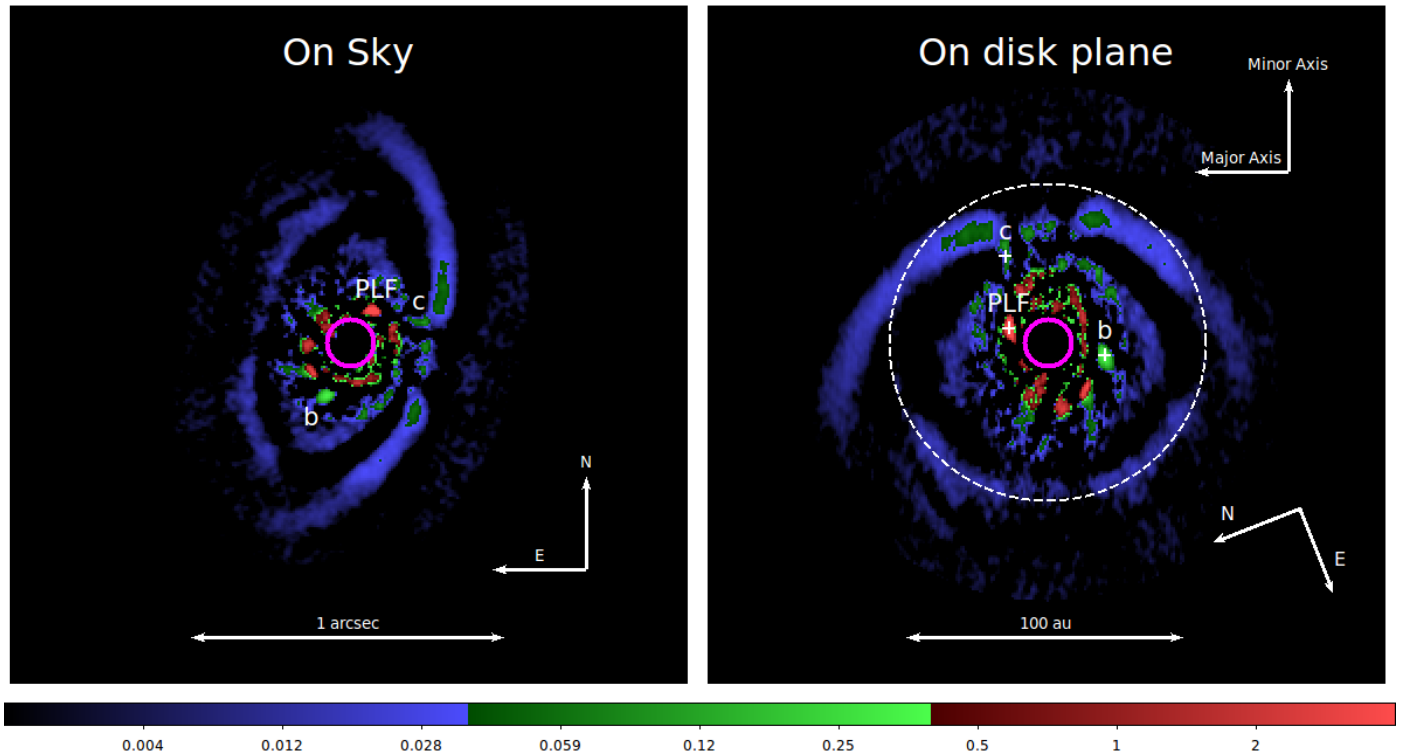


Fig. 2. *Left:* IFS image obtained combining the images from all the epochs as described in the text. The positions of the three objects known around PDS 70 are tagged. *Right:* Deprojected image of the PDS 70 system as seen on the disk plane. The dashed circle indicate the position of the inner edge of the outer ring of the disk as found with ALMA by [Keppler et al. \(2019\)](#) to show that it corresponds to the partially imaged ring in the SPHERE images.

played in panel *a* of Figure 5 where the green squares are the photometric values obtained from IFS while the violet squares are obtained from IRDIS. The extracted spectrum displays very red colors providing further confirmation of the planetary nature of this object as also confirmed by comparing it to the solid cyan line that represents the extracted spectrum for a speckle from the same dataset, obtained before applying the speckle subtraction procedure at a separation comparable to that of the companion. Using these results we can also derive the photometric results in different spectral bands that are listed for the two epochs in which it is visible in Table 2.

3.3. PLF

Like for the case of PDS 70 c we extracted astrometric and photometric values for the PLF following the same methods. In this case however, we were able, using only the IFS data, to obtain the astrometric values in all four observing epochs exploited for this work. This supports the fact that the signal is not an artefact from the data processing. The astrometric values for this object are listed in Table 3 while its relative positions are compared to those expected for a background object in the bottom right panel of Figure 3. Also in this case the analysis confirms that this source is gravitationally bound to PDS 70. As done for PDS 70 c, the spectrum extraction was performed both using the negative planet procedure, the ANDROMEDA package and the PACO package leading to comparable results for the IFS data. Due to the fact that it was possible to detect this object only marginally in IRDIS data, it was not possible to extract any reliable photometry with the first two methods while this was possible using the PACO package. We then use these values for the spectrum of this object that is shown in panel *b* of Figure 5. PACO could lead to an underestimation of the error bars with respect to other methods due to a different way to calculate them as can be seen in Figure 5. While this can lead to some inconsistency, we included its results to give a more complete view of the characteristics of this object. It appears blue and very similar to the stellar spectrum at these wavelengths as demonstrated by the blue solid line plotted in the same Figure obtained using, following Müller et al. (2018), a calibrated spectrum of PDS 70 from the SpeX spectrograph. To be able to overplot the two spectra, we have normalized the stellar spectrum to the flux of the companion. This is a strong hint that the PLF spectrum is probably mainly due to the reflection of the stellar light from dust.

3.4. Structure of the PDS 70 planetary system

To constrain the fundamental orbital parameters for the three candidate companions, we run a Monte-Carlo simulation using the Thiele-Innes formalism (Binnendijk 1960) as devised in Desidera et al. (2011), Zurlo et al. (2013) and Zurlo et al. (2018) and adopting the convention by Heintz (2000). The simulation generates 5×10^7 orbits and rejects all those that do not fit the astrometric data. With this aim, we assumed for the star a mass of $0.76 M_{\odot}$ and a distance of 113.43 pc. Moreover, we assumed a circular orbit for each companion. In this way, we found for PDS 70 b an orbital radius of $22.7^{+2.0}_{-0.5}$ au, a period of $123.5^{+9.8}_{-4.9}$ years and an orbit inclination of $39.7^{+5.4}_{-2.8}$, in good agreement with what was found by Müller et al. (2018). Due to the paucity of astrometric points, the results for PDS 70 c are less defined than for PDS 70 b. We obtained a radius of $30.2^{+2.0}_{-2.4}$ au, a period of $191.5^{+15.8}_{-31.5}$ years and an inclination of $40.8^{+30.3}_{-14.0}$ degrees. Finally, we obtained for the PLF orbit a ra-

dius of $13.5^{+0.3}_{-0.2}$ au, a period of $56.3^{+2.0}_{-1.1}$ years and an inclination of $46.5^{+2.3}_{-10.4}$ degrees.

We combined these results together with those for the disk (Keppler et al. 2019) to generate a model for the structure of the planetary system around PDS 70. The values of the estimated separation expressed in au both for the companion and for the part of the disk, together with the Keplerian rotation period, the companion masses as calculated in Section 4.1 and 4.2 and with the ratio between the object period and the period of PDS 70 b are listed in Table 4. From these data we can see that, if the PLF is a planet, its orbit is in 2:1 resonance with the outer edge of the inner disk while a 3:2 resonance can be found between the orbits of PDS 70 b and PDS 70 c.

4. Discussion

4.1. On the nature of PDS 70 c

The projected position of PDS 70 c is very close to the bright West part of the disk. This rises the concern that the spectrum shown in the panel *a* of Figure 5 can be contaminated and hence biased by the disk signal. To check this possibility we have extracted the spectrum of the disk (using both IFS and IRDIS photometry) at ten different positions of the disk as indicated by the red circles in panel *f* of Figure 1. Due to the fact that we are extracting the photometry of an extended structure, it is not possible to use the negative planet procedure as we did for PDS 70 c. We then applied an aperture photometry method, using an aperture with a radius of three pixels. We then made a median of these spectra to compare the resulting spectrum with that extracted for the companion. The result of this procedure is displayed in panel *c* of Figure 5 where the spectra of the companion and of the disk are represented by the green and the orange line, respectively. We can see that, while at short wavelengths the two spectra are very similar, they diverge starting from the IFS H-band and are very different in the IRDIS K-band. As a further test, to check if the disk can influence the companion spectrum extraction, we have injected in our datacube a simulated planet in a position similar to PDS 70 c with respect to the disk. The position of the simulated companion is indicated by a green circle in panel *f* of Figure 1. The input spectrum of the simulated planet was obtained scaling the spectrum of PDS 70 b. After that, we extracted the companion spectrum using the same method as used for PDS 70 c. The result of this procedure is displayed in panel *d* of Figure 5 where the input spectrum is represented by the blue line while the orange squares are the extracted photometric values. From these results we can conclude that we are able to faithfully reproduce the input spectrum of the simulated planet. Given the similarity of this case with that of PDS 70 c, we conclude that the results of these tests support the hypothesis that the signal comes from a point-like source not related to the circumstellar disk and is of planetary origin. However, it is important to stress that in the Y- and J-spectral bands, the signals from the companion and from the disk are both very low and virtually indistinguishable so that, while we will use them to define the physical characteristics of the companion, the final results should be regarded with caution providing just upper limits.

Moreover, we can check the reliability of the extracted spectrum by comparing it to the spectrum of PDS 70 b. In panel *a* of Figure 5 we overplot the extracted spectrum for PDS 70 b, rescaled to be able to compare it with that of PDS 70 c. For this purpose, we used a calculated rescaling factor of 0.18. The two spectra are very similar indicating that both these objects possess red and dusty atmospheres and/or circumplanetary disks.

Table 2. Astrometric and photometric (absolute magnitudes) results obtained for PDS 70 c.

Date	ΔRA (")	ΔDec (")	ρ (")	PA	M_J	M_H	M_{K1}	M_{K2}
2018-02-24	-0.205 ± 0.013	0.041 ± 0.006	0.209 ± 0.013	281.2 ± 0.5	17.45 ± 0.33	14.89 ± 0.87	12.69 ± 0.17	12.53 ± 0.19
2019-03-06	-0.222 ± 0.008	0.039 ± 0.004	0.225 ± 0.008	279.9 ± 0.5	17.00 ± 0.40	14.90 ± 0.90	12.49 ± 0.11	12.25 ± 0.14

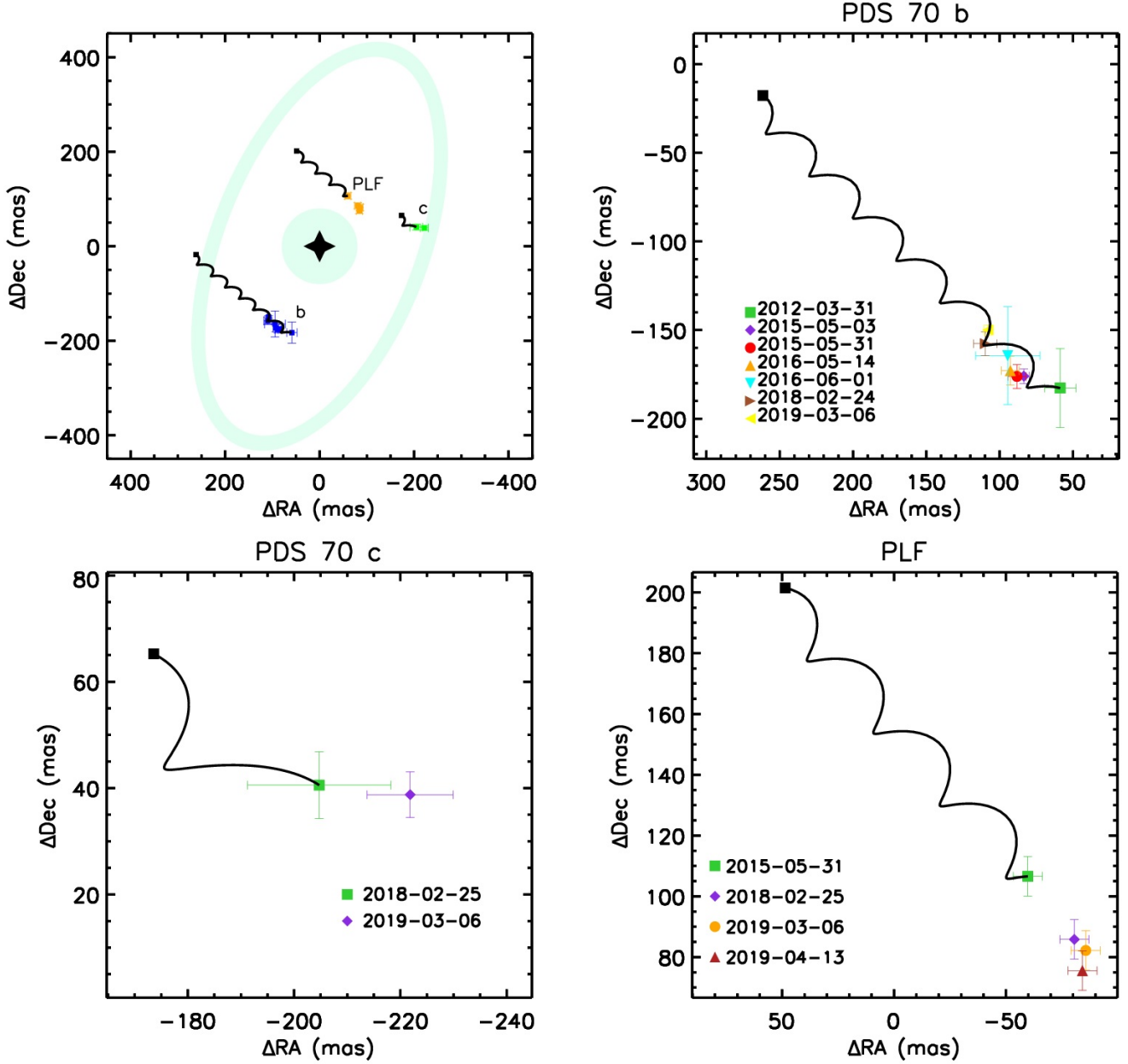


Fig. 3. *Top left panel:* Relative astrometric positions of the three proposed companions of PDS 70 with respect to the host star, represented by the black star symbol. The positions of the inner and of the outer disk are also displayed through light cyan areas. The solid black lines represent the expected course of the companion if it were a background object. The length differences between the black lines for different objects are due to the different temporal coverage that we have for each of them. The black squares at the end of the line represent the expected position at the epoch of the last observation in this case. For clarity, we represent a zoomed version of the areas around each companion in the *Top right panel* for PDS 70 b, in the *Bottom left panel* for PDS 70 c and in the *Bottom right panel* for the PLF.

We have estimated the mass of PDS 70 c by comparing the photometric values reported in Table 2 with the AMES-DUSTY models (Allard et al. 2001) assuming the system age of 5.4 ± 1.0 Myr. The error bars on the age could be underestimated for late-type Scorpius-Centaurus stars due to the fact that the effects of the magnetic activity are not taken properly into

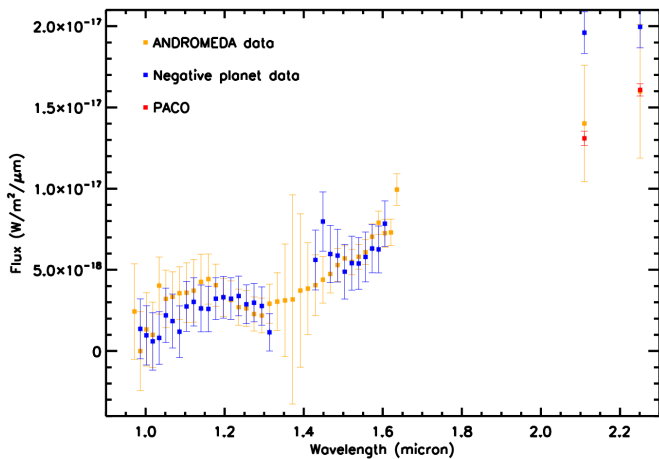
account as found by Asensio-Torres et al. (2019) for the case of HIP 79124. The results of this procedure are listed in Table 5. The error bars are obtained taking into account the uncertainties both on the distance of the system, on the age of the system and on the extracted photometry. The mass values estimated are in good agreement between the different wavelength bands and

Table 3. Astrometric results obtained for the PLF.

Date	ΔRA (")	ΔDec (")	ρ (")	PA
2015-05-31	-0.060 ± 0.007	0.107 ± 0.007	0.122 ± 0.007	330.7 ± 0.5
2018-02-24	-0.081 ± 0.004	0.086 ± 0.004	0.118 ± 0.004	316.8 ± 0.5
2019-03-06	-0.086 ± 0.004	0.082 ± 0.004	0.119 ± 0.004	313.8 ± 0.5
2019-04-13	-0.084 ± 0.004	0.076 ± 0.004	0.113 ± 0.004	311.7 ± 0.5

Table 4. Structure of the PDS 70 planetary system. In second columns we list the masses of the of the companions including also that of the PLF as it was point source. In third column we list the separations both for companions and for inner and outer disk. In fourth column we listed the most probable periods (for the disk part we assume a keplerian motion). Finally, in fifth column we list the ratio between the period of the companion/disk and the period of PDS 70 b to evidence possible orbital resonance.

Object	Mass (M_{Jup})	Separation (au)	Period (year)	$P_{\text{obj}}/P_{\text{PDS 70 b}}$
Inner disk	//	9.1	31.5	0.26
PLF	0.05-0.28	$13.5^{+0.3}_{-0.2}$	$56.3^{+2.0}_{-1.1}$	0.46
PDS 70 b	5.0-9.0	$22.7^{+2.0}_{-0.5}$	$123.5^{+9.8}_{-4.9}$	1.00
PDS 70 c	4.4 ± 1.1	$30.2^{+2.0}_{-2.4}$	$191.5^{+15.8}_{-31.5}$	1.55
Outer disk (ring 1)	//	60.0	533.1	4.32
Outer disk (ring 2)	//	74.0	730.2	5.91


Fig. 4. Comparison between the PDS 70 c spectrum extracted using the negative planet method (blue squares), ANDROMEDA (orange squares) and PACO (magenta squares). For the latter method, we plot only the two values obtained from IRDIS data. Due to fact that in the region around $1.4 \mu\text{m}$ the negative planet method gives negative values, those points are not shown in the plot.

point toward a planetary object with a mass just above $4 M_{\text{Jup}}$. We notice here that these determinations assume that the observed magnitudes are due to the planetary atmosphere and that the age of the companion is the same as the age of the star.

We have also used the extracted photometry to build two color-magnitude diagrams. They were created following the procedure described in Bonnefoy et al. (2018) and they are displayed in the left (J-H versus J) and in the right (J-K1 versus J) panels of Figure 6. The positions of PDS 70 c in these two diagrams demonstrate that it is an extremely red object even with respect to PDS 70 b that is also shown for comparison. The larger redness of PDS 70 c is also confirmed by the comparison of the spectra of these two objects in panel a of Figure 5 where the K-band values for PDS 70 b are lower than those for PDS 70 c when the rescaling factor is taken into account. We overplot on both color-magnitude diagrams the reddening vectors due to interstellar extinction (Draine 2003) and by $0.5 \mu\text{m}$ forsterite grains using the optical constants by Scott & Duley (1996). These grains are

Table 5. Mass estimation for PDS 70 c obtained from the SPHERE photometry and from the AMES-DUSTY models.

Spectral band	Mass (M_{Jup})
J	4.35 ± 0.45
H	4.21 ± 0.78
K1	4.31 ± 0.45
K2	4.02 ± 0.41

proposed to explain the red colors of dusty and variable L dwarfs (Marocco et al. 2014; Bonnefoy et al. 2016). Both PDS 70 b and PDS 70 c appear to be shifted along these vectors with respect to the sequence of field dwarfs. These results are compatible with both these objects being low mass objects strongly reddened by the presence of dust.

With the aim to characterize this companion we fitted its extracted spectrum with Exo-REM models taking into account the presence of thick clouds (Baudino et al. 2015, 2017; Charnay et al. 2018). The atmospheric models were calculated on a grid with T_{eff} varying between 300 and 2000 K with steps of 50 K and with $\log g$ with values between 3.0 and 6.0 dex with steps of 0.1 dex. On the other hand, we verified that the metallicity did not influence the final result so that we assumed solar metallicity (we explored also the 0.3x and 3x solar metallicity). As can be seen in the panel a of Figure 5, where the orange solid line represents the best fit Exo-REM model, the fit is good along all the extracted spectrum. Best fits are found for models with low gravity ($\log g \sim 3.0 - 3.5$) and for T_{eff} in the range 800-1100 K. The masses range between 0.5 and $4 M_{\text{Jup}}$ and are just partially in agreement with what we found previously with AMES-DUSTY models. The best fit model represented in Figure 5 has $T_{\text{eff}}=900$ K, $\log g=3.1$ dex, $R=1.95 R_{\text{Jup}}$ and $M=1.93 M_{\text{Jup}}$.

To confirm these results we performed a similar procedure using the BT-Settl models (Allard 2014) with T_{eff} varying between 700 and 2500 K with a step of 100 K and a $\log g$ between 3.0 and 5.5 dex with a step of 0.5 dex. Also in this case we considered only solar metallicity. We obtained good fit for T_{eff} between 800 and 1100 K and for $\log g=3.5$. We can then conclude that the two models are in good agreement as also confirmed by the best fit model also displayed in Figure 5 with the green

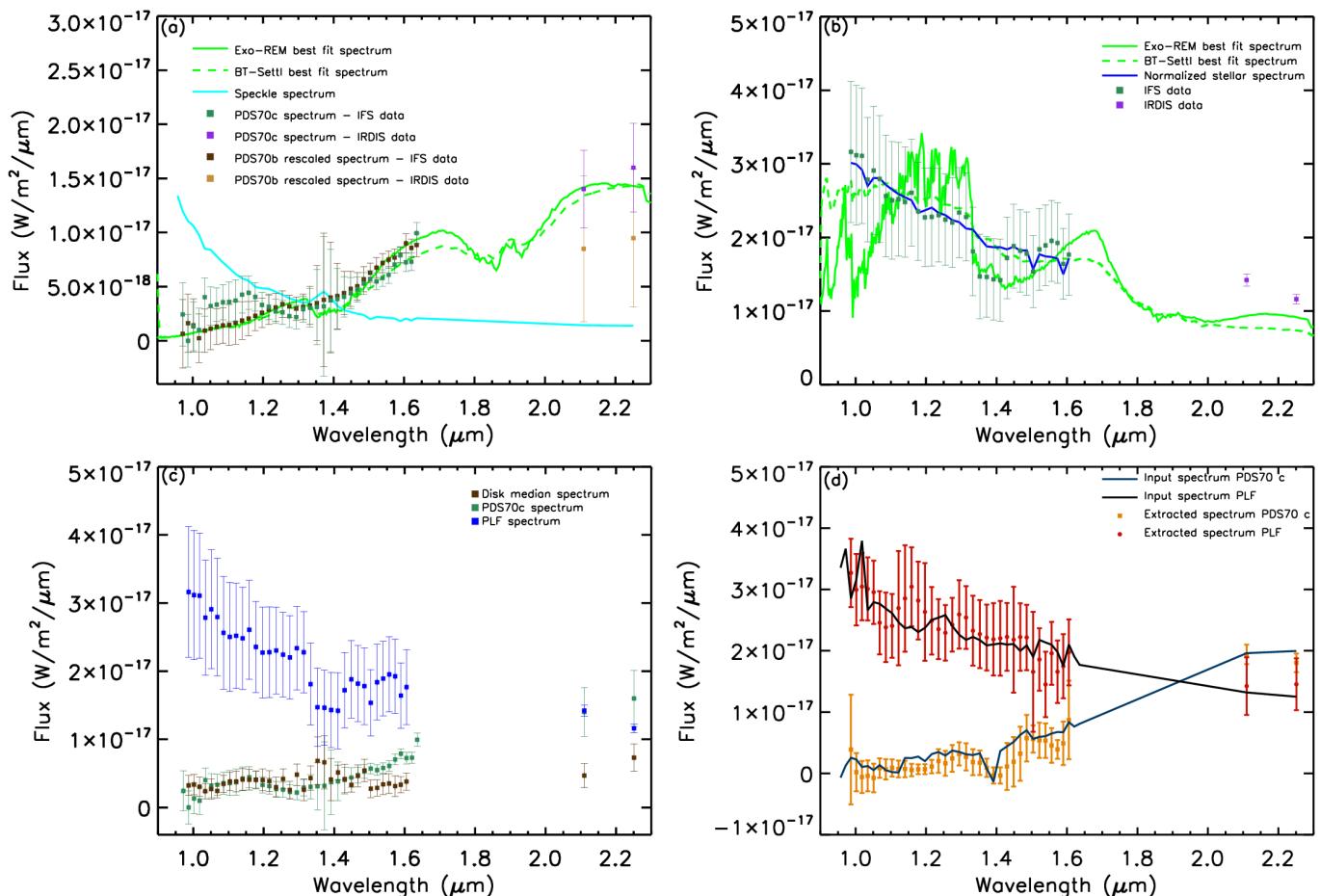


Fig. 5. *Panel (a):* Extracted spectrum of PDS 70 c both with IFS (green squares) and IRDIS (violet squares). The green solid line represents the best fit Exo-REM model while the green dashed line is the best fit BT-Settl model. The parameters of the best-fit models are listed in the text (Section 4.1). The cyan solid line represents an extracted spectrum for a speckle of the same dataset. We also overplot the rescaled spectrum of PDS 70 b both for IFS (deep brown squares) and IRDIS (light brown squares). *Panel (b):* Same as *panel (a)* but for the PLF spectrum. We plot as a blue solid line the stellar spectrum normalized to the companion flux to demonstrate their similarity. *Panel (c):* Comparison between the extracted spectra for PDS 70 c and the PLF and the spectrum obtained making a median of ten different positions of the disk. *Panel (d):* Comparison between the injected spectrum of a simulated planet in a position similar to PDS 70 c with respect to the disk (deep blue solid line) and the relative extracted spectrum (orange squares). The same thing is done for a simulated planet at a similar position of the PLF (black solid line and red circles).

solid line. In this case the best fit is obtained for $T_{\text{eff}}=900$ K and $\log g=3.5$ dex.

The recently suggested presence of a circumplanetary disk (Isella et al. 2019) may cause selective absorption, reddening the object spectrum. This would explain, at least partially, the very red colors that we have found in our work and it would confirm recent results on the detectability of circumplanetary disks from hydrodynamic simulation (Szulágyi & Garufi 2019; Szulágyi et al. 2019). The effects of the circumplanetary disk should be carefully modeled because they could partially modify the fitting results that we have found in this Section but this analysis is beyond the scope of the present work. The results presented here should be regarded as a first estimation of the companion physical characteristics.

4.2. On the nature of the PLF

The fact that the PLF is recoverable in at least four different epochs is a clear evidence that it is a real astrophysical signal. Moreover, its astrometric positions in the four epochs clearly exclude the possibility that it is a background object and they

are in agreement with a Keplerian circular motion around the star. However, as discussed above, its spectrum is blue and very similar to the stellar spectrum. To confirm the reliability of the extracted spectrum we have extracted the spectrum of a simulated planet with a spectrum similar to that of the PLF and in a position similar to that of the PLF following the same method used for PDS 70 c and described in Section 4.1. The results are displayed in panel *d* of Figure 5 where the input spectrum is represented with a black solid line while the extracted spectrum is represented by red circles. Like for the case of PDS 70 c, we are able to faithfully reproduce the input spectrum. A blue spectrum is a strong hint that the emission is due to stellar light reflected likely by dust. Its measured separation of ~ 13.5 au is not consistent with the radius of the inner disk of about 9 au (Keppler et al. 2019) corresponding to an angular radius of $\sim 0.08''$ that should be completely behind the SPHERE coronagraph. We have however to consider that the inner disk size is obtained in the sub-millimeter continuum that probes the distribution of large dust grains. We cannot exclude that the distribution of small dust grains, probed by scattered light, could be different.

A hint favoring this latter view comes from the analysis of the IFS off-axis PSF images obtained just before and just after

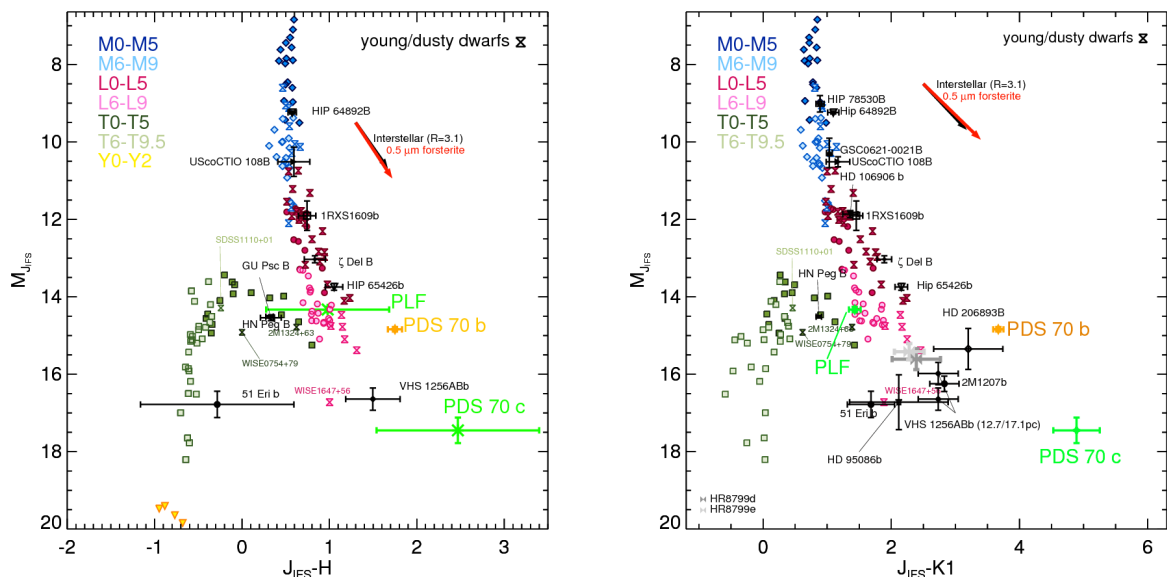


Fig. 6. *Left:* Positions of PDS 70 c and of the PLF (green stars) in the J-H versus J color-magnitude diagram. The position of PDS 70 b (orange star) is also indicated as comparison. We also display the positions of field dwarfs indicated with different symbols according to the spectral type and of some low mass companions. Moreover, we overplot the reddening vector computed for the synthetic extinction curve using $R_V=3.1$ (black arrow) and the reddening vector due to the extinction curve of forsterite (red arrow). *Right:* Same as left panel for the J-K1 versus J diagram.

the coronagraphic observations as explained in Section 2. Exploiting the field rotation between these two images we can perform a subtraction to eliminate static speckles on these images. This procedure is performed separately in Y, J and H band. In normal conditions, the limiting contrast that we can obtain in the region between 60 and 110 mas from the star is of the order of 6-7 magnitudes. However, the very good conditions obtained during the night of 2018-02-24 allowed us to reach a contrast of the order of 10 magnitudes at the same separations. In Figure 7 we present the results of this procedure on the H-band data. In this image there is a possible detection of the inner ring of the disk. We draw an ellipse (red solid line) adopting the inclination and the PA of 49.7° and 158.6° , respectively. Moreover, its semi-major axis is of 128 mas. This ellipse is roughly passing through the position of the PLF. While this is a differential image, it is not an ADI image because it uses only two frames. For this reason, we expect that the region on the near side of the disk around the semi-minor axis (corresponding to the forward scattering of the disk) is the brightest, as indeed it is. The contrast value in the position of the PLF is about 9.5 mag, consistent with the photometry obtained on the coronagraphic image. These results therefore appear consistent with scattered light from the inner disk. In any case, it is important to stress that while these results are suggestive of the presence of the inner disk, they cannot be considered as a clear detection due to the very low SNR of the signal and to the presence of other faint structures in the FOV. However, we would like to note that this structure has the orientation expected for the inner disk and that a comparison with other observations taken in similar optimal conditions as those for PDS 70 does not reveal similar structures. If the interpretation of this image as the inner disk is correct, then the PLF is a part of the disk itself.

The positions of the PLF in both the color-magnitude diagrams displayed in Figure 6 are compatible with those of field dwarfs with spectral type L6-L9. Moreover, we followed the same procedure used for PDS 70 c fitting its extracted spectrum both with Exo-REM and BT-Settl atmospheric models. The best fit spectra from this procedure are shown in the panel *b* of Fig-

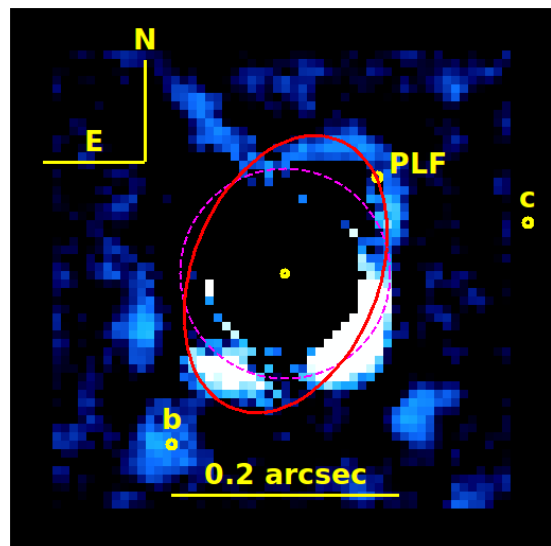


Fig. 7. Image resulting from the off-axis PSF subtraction described in the text. The red ellipse represent the estimated position for the inner disk. The dashed magenta circle represents the position of the coronagraph. The yellow circles are used to tag the positions of the point sources.

ure 5. Exo-REM models point toward a small star with a mass of $\sim 100 M_{\text{Jup}}$ and T_{eff} between 1400 and 2000 K. The value of $\log g$ should be of the order of 6. On the other hand, BT-Settl models favor a much higher T_{eff} of the order of 3000 K while the surface gravity is unconstrained. These results are generally quite inconsistent with each other. Moreover, a small star at such a short separation from the star would probably have a strong effect on the stability of the system generating a strong radial velocity signal that has never been observed. Furthermore, a stellar object would be much brighter, of the order of 6-9 magnitudes, than what we found for the PLF. These considerations seem to confirm our previous conclusions that we are not actually looking to

the photosphere of a sub-stellar object but rather to stellar light reflected by dust. If so, the inconsistent results obtained when attempting to fit this spectrum with the atmospheric models are naturally explained.

The results above seem to demonstrate that the PLF is actually a part of the inner disk, which is supported by the fact that no H_α emission associated to the PLF has been detected so far. Still, we cannot completely exclude the possibility that it is actually a point source that should be however embedded into the disk. If this were the case, we can consider two alternative possibilities. The first one is that it is a transient blob of dust that is destined to break apart in a timescale of few thousands of years similar to blobs B and C seen in the disk of HD 169142 (Ligi et al. 2018; Gratton et al. 2019) and, for this reason, the probability to observe such a transient phenomenon is generally low unless these structures are generated frequently. The second possibility is that we are observing a dust envelope around a forming planet. This case would be very similar to what was suggested for blob D around HD 169142 with SPHERE data (Gratton et al. 2019). Given the spectrum, we should not see any photospheric emission from the planet due to the fact that the envelope is optically thick or the light reflected by the dust dominates over the photospheric emission. If this is the case, we can estimate a guess for the mass of the planet considering its contrast and assuming that the dust is filling the Hill radius around the planet. The contrast of the object is given by:

$$C = A \cdot \frac{\pi r_H^2}{4\pi a^2} \cdot \frac{1}{k} \quad (1)$$

where A is the albedo, r_H is the Hill radius, a is the separation of the object from the star and k is a multiplicative factor depending on the geometry of the system. Considering the inclination of the system, we found that a reasonable value of k is 1.38. The Hill radius is given by:

$$r_H = a \cdot \sqrt[3]{\frac{m_p}{3M_\star}} \quad (2)$$

where m_p is the mass of the planet and M_\star is the stellar mass. Joining equation 1 and equation 2 we obtain an expression for the planetary mass:

$$m_p = 3M_\star \cdot \left(\frac{4Ck}{A}\right)^{\frac{3}{2}} \quad (3)$$

The value of the contrast is 7.27×10^{-5} as obtained from a median of the contrast values on each IFS spectral channel. We then assume for the star a mass of $0.76 M_\odot$ (Müller et al. 2018) and an albedo (A) of 0.5. We then obtain a value for the mass of the planet of $5.2 \times 10^{-5} M_\odot$ corresponding to $\sim 17.3 M_\oplus$. This is of course a lower limit valid in the case that the circumplanetary material is filling the Hill radius. In a more evolved case, in which the circumplanetary material is filling 1/3 of the Hill radius as proposed e.g. by Ayliffe & Bate (2009) the mass of the planetary objects would be of the order of $90 M_\oplus$. If this were true, in the PDS 70 system we would have planets at a different evolution stage. Indeed, while PDS 70 b and PDS 70 c are accreting material probably from a circumplanetary disk, this object would still be in its formation phase. This would probably recall what happened in our solar system where the gaseous giant planets formed before the inner planets.

4.3. Stability of the system

To test the stability of the system, we performed a few N-body simulations with the integrator SWIFT HJS (Beust 2003). This method is optimized to check what happens in the disk cavity because it does not take into account the gas component of the disk. The main assumption of this model is that there is no major influence of the circumstellar disk on the planetary orbits. The critical parameters that control the stability of the system are the semi-major axes ratios and the eccentricities. To this aim we used the values listed in Table 4 for PDS 70 b and PDS 70 c. Given the current observational constraints, the stability of the system is not guaranteed if the masses are as large as obtained assuming an age of 5.4 Myr for the planets. The companions are not sufficiently separated to avoid significant gravitational interactions in a timescale of the order of few thousands of years that is not compatible with the age of the system. The chaotic zones around each one of them can be computed from Morrison & Malhotra (2015). The case of coplanar circular orbits at the observed separations is represented in Figure 8. The separation of PDS 70 c falls within the chaotic zone of PDS 70 b, indicating strong dynamical perturbations. If we assume that the PLF is actually a massive body with the mass determined in Section 4.2 the results on the stability do not change due to its proposed low mass. As shown in Figure 9, it is outside the chaotic zones of its neighbor, but not by far, so that eccentricity fluctuations could eventually destabilize it. Assuming circular orbits for all three companions, to limit dynamical perturbations, still leads to a destabilization within a few thousand years. Given the age of the system, this is not likely to be the case. The companion separations are compatible with a configuration of mean-motion-resonances 1:2:3, and marginally, with a 1:2:4 configuration. This system would be thus very similar to the 4-planet system HR 8799 (Marois et al. 2008), who also appeared unstable and that is potentially in resonant configuration (Goździewski & Migaszewski 2009) or stable at the condition that the planets are coplanar or in mean-motion resonance (Konopacky et al. 2016; Wang et al. 2018). A comprehensive study of the dynamics of PDS 70 would need to take into account both the mean-motion resonances and the interactions with the disk, which could act as a stabilizing factor by damping the eccentricities. On the other hand, at first order, both the inner disk and outer disk edges are compatible with the positions of the bodies. Furthermore, the disk is susceptible to be affected by the gravitational influence of the PLF, if a real object, for the inner disk and PDS 70 c for the outer disk, through secular dynamics and mean-motion resonances. This could cause local depletions in some part of the disk or asymmetries. However, to study in more detail these effects we would need hydrodynamical simulations of the disk evolution and this is outside the scope of this paper.

We would like to stress that these results have been achieved assuming for PDS 70 b and PDS 70 c masses obtained using the age system of 5.4 Myr. Anyhow, given that these two objects are still in formation, their ages could be younger and, as a consequence, also their masses would be lower changing the results about the stability of the system. Indeed, stability simulations performed with lower masses for PDS 70 b and PDS 70 c ($2 M_{\text{Jup}}$), would allow to reach the stability also when we change only the mass of one of the two planets if the separation of PDS 70 c is 31 au or larger. The orbital parameters of this object were determined thanks to only two astrometric points so that further astrometric measures will be needed to better constrain them. If we instead consider for these two objects the two masses given in Table 4, we would need for PDS 70 c a separation larger

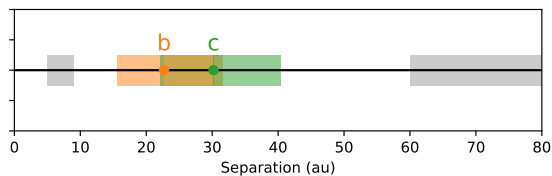


Fig. 8. Positions of the chaotic zones for PDS 70 b and PDS 70 c. The grey zones represent the disk positions.

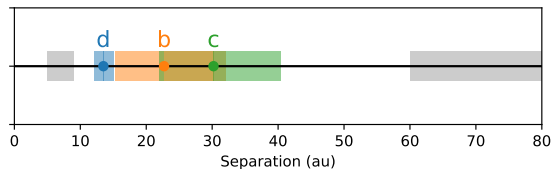


Fig. 9. Similar to Figure 8 but including also the PLF (here tagged as d) considering it as a real companion.

than 35-36 au or alternatively a separation for PDS 70 b smaller than 20 au, obviously just in the case that the PLF is not a companion.

The system stability could be enhanced if the two planets are in mean motion resonance as found recently by [Bae et al. \(2019\)](#) using two-dimensional hydrodynamic simulations.

5. Conclusions

In this work we presented the SPHERE observations aimed to verify the presence of a second planetary companion around the star PDS 70. Our results confirm the presence of PDS 70 c both in J- and H-spectral bands with IFS and in K band with IRDIS despite the flux is very low at shorter wavelengths. We retrieved astrometric positions for this object in two different epochs and this allowed us to support that this object is actually gravitationally bound to the star. The extracted photometry hints toward a very red spectrum that is very probably due to the presence of dust in the planetary atmosphere and/or to the presence of a circumplanetary disk as proposed through recent ALMA observations. We were able to obtain a good fit with atmospheric models and our results clearly tend toward a low gravity (3.0-3.5 dex) and low temperature (~ 900 K) object using both the Exo-REM and the BT-Settl models. The mass evaluated through these models is just in marginal agreement with that obtained through the AMES-DUSTY evolutionary models, but we can however conclude that it should be less than $\sim 5 M_{\text{Jup}}$. Finally, the probable presence of dust in the atmosphere of this object and/or of a circum-planetary disk is also confirmed by its position in the color-magnitude diagram that is even more extreme than that of PDS 70 b.

In addition to the results on PDS 70 c, the analysis of archival and new data allowed us to detect the presence of a possible third point source in this planetary system (PLF). This object is located at small separation ($\sim 0.12''$) from the star and it is clearly detected only in the IFS data while it is barely visible in the IRDIS data. In any case, it was possible to detect it in four different epochs with a time range of about three years allowing to exclude that it can simply be a feature due to the speckle subtraction method. Also, these astrometric positions allowed us to confirm that this object is actually gravitationally bound to the star. However, the extracted spectrum for this object is blue hinting toward stellar light reflected from dust. This result seems

to demonstrate that this object is actually a feature due to the inner disk. This is further confirmed by an analysis of the non-coronagraphic PSF data from which we were able to find hints of the presence of the inner disk at the separation of this feature. We cannot, however, definitely exclude the possibility that it is a separate object and in this case we considered two possible solutions. The first one is that it is just a transient blob of dust while in the second case it could be a forming planet completely embedded in a dust envelope. In the first case we should assume that the generation of such blobs is frequent through some not well defined mechanism. In the second case we would not be able to see any light from the planetary photosphere and this would justify the extracted spectrum. If this is the real nature of this object we can evaluate its mass being less than $90 M_{\oplus}$. It is not possible with the present data to draw a definitive conclusion about the nature of this object. To be able to obtain a definitive answer on this point, interferometric data with instruments like e.g. GRAVITY ([Gravity Collaboration et al. 2017](#)) of the inner part of this system would be mandatory.

We also presented a possible solution for the general structure of the planetary system assuming for the moment circular orbits for all the candidate companions. Moreover, the simulations were performed both excluding the PLF and both considering it. In both cases, the proposed configuration does not seem to be stable on long periods if the masses are as large as obtained assuming that the age of the companions is the same as for the star. In any case, as noted above, lower masses of the companions due to a lower age of the forming planets with respect to the system would allow us to obtain a stable configuration. Clearly long period astrometric follow-up observations will be needed to further constrain the orbits of the proposed objects and to draw more precise conclusion about the structure and the stability of the PDS 70 system.

From our data we can conclude that PDS 70 hosts the second multiple planetary system found with the direct imaging technique and the first one at a very young age, where the planets are still accreting.

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¹INAF-Osservatorio Astronomico di Padova, Vicolo dell'Osservatorio 5, Padova, Italy, 35122-I

²Max-Planck-Institut für Astronomie, Königstuhl 17, 69117, Heidelberg, Germany

³Univ. Grenoble Alpes, CNRS, IPAG, 38000 Grenoble, France

⁴LESIA, Observatoire de Paris, PSL Research University, CNRS, Sorbonne Universités, UPMC Univ. Paris 06, Univ. Paris Diderot, Sorbonne Paris Cité, 5 Place Jules Janssen, 92195 Meudon, France

⁵Univ. Lyon, Univ. Lyon 1, ENS de Lyon, CNRS, CRAL UMR 5574, 69230 Saint-Genis-Laval, France

⁶Aix Marseille Univ., CNRS, CNES, LAM, Marseille, France

⁷Université de Lyon, UJM-Saint-Etienne, CNRS, Institut d'Optique Graduate School, Laboratoire Hubert Curien UMR 5516, F-42023, Saint-Etienne, France

⁸Department of Terrestrial Magnetism, Carnegie Institution for Science, 5241 Broad Branch Road, NW, Washington, DC 20015, USA

⁹Unidad Mixta Internacional Franco-Chilena de Astronomía (CNRS, UMI 3386), Departamento de Astronomía, Universidad de Chile, Camino El Observatorio 1515, Las Condes, Santiago, Chile

¹⁰Leiden Observatory, Leiden University, PO Box 9513, 2300 RA Leiden, The Netherlands

¹¹Department of Astronomy, University of Michigan, 1085 S. University Ave, Ann Arbor, MI 48109-1107, USA

¹²European Southern Observatory (ESO), Alonso de Cordova 3107, Vitacura, Casilla 19001, Santiago, Chile

¹³SPGroup, ELEN/ICTEAM, UCLouvain, B-1348 Louvain-la-Neuve, Belgium

¹⁴Núcleo de Astronomía, Facultad de Ingeniería y Ciencias, Universidad Diego Portales, Av. Ejercito 441, Santiago, Chile

¹⁵Escuela de Ingeniería Industrial, Facultad de Ingeniería y Ciencias, Universidad Diego Portales, Av. Ejercito 441, Santiago, Chile

¹⁶INAF – Osservatorio Astronomico di Roma, via di Frascati 33, 00078 Monte Porzio Catone, Italy

¹⁷Department of Physics, University of Oxford, Oxford OX1 3PU, UK

¹⁸INAF - Osservatorio Astronomico di Capodimonte, Salita Moiariello 16, 80131 Napoli, Italy

¹⁹Geneva Observatory, University of Geneva, Chemin des Maillettes 51, 1290 Versoix, Switzerland

²⁰Department of Astronomy, Stockholm University, Stockholm, Sweden

²¹Université Côte d'Azur, OCA, CNRS, Lagrange, France

²²INAF – Osservatorio Astronomico di Brera, Via E. Bianchi 46, I-23807 Merate, Italy

²³STAR Institute, Université de Liège, Allée du Six Août 19c, B-4000, Liège, Belgium

²⁴Instituto de Física y Astronomía, Facultad de Ciencias, Universidad de Valparaíso, Av. Gran Bretaña 1111, Valparaíso, Chile

²⁵Núcleo Milenio Formación Planetaria – NPF, Universidad de Valparaíso, Av. Gran Bretaña 1111, Valparaíso, Chile

²⁶Monash Centre for Astrophysics (MoCA) and School of Physics and Astronomy, Monash University, Clayton Vic 3800, Australia

²⁷DOTA, ONERA, Université Paris Saclay, F-91123, Palaiseau France

²⁸Hamburger Sternwarte, Gojenbergsweg 112, D-21029 Hamburg, Germany